

Planck Early Results XI: Calibration of the local galaxy cluster Sunyaev-Zeldovich scaling relations

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ABSTRACT

We present precise Sunyaev-Zeldovich (SZ) effect measurements in the direction of 62 nearby galaxy clusters ($z < 0.5$) detected at high signal-to-noise in the first *Planck* all-sky data set. The sample spans approximately a decade in total mass, $2 \times 10^{14} M_{\odot} < M_{500} < 2 \times 10^{15} M_{\odot}$, where M_{500} is the mass corresponding to a total density contrast of 500. Combining these high quality *Planck* measurements with deep *XMM-Newton* X-ray data, we investigate the relations between $D_A^2 Y_{500}$, the integrated Compton parameter due to the SZ effect, and the X-ray-derived gas mass $M_{g,500}$, temperature T_X , luminosity $L_{X,500}$, SZ signal analogue $Y_{X,500} = M_{g,500} \times T_X$, and total mass M_{500} . After correction for the effect of selection bias on the scaling relations, we find results that are in excellent agreement with both X-ray predictions and recently-published ground-based data derived from smaller samples. The present data yield an exceptionally robust, high-quality local reference, and illustrate *Planck*'s unique capabilities for all-sky statistical studies of galaxy clusters.

Key words. Cosmology: observations, Galaxies: cluster: general, Galaxies: clusters: intracluster medium, Cosmic background radiation, X-rays: galaxies: clusters, Planck satellite

1. Introduction

The X-ray emitting gas in galaxy clusters induces inverse Compton scattering of Cosmic Microwave Background (CMB) photons, shifting their frequency distribution towards higher energies. First discussed in 1972 by [Sunyaev & Zeldovich](#), the scattering produces a characteristic distortion of the CMB spectrum in the direction of a cluster known as the thermal Sunyaev-Zeldovich (SZ) effect. It is directly proportional to the Compton parameter y , a measure of the thermal electron pressure of the intracluster medium (ICM) gas along the line of sight: $y =$

$(\sigma_T/m_e c^2) \int P dl$. Here $P \propto n_e T$ is the ICM thermal electron pressure, where n_e is the density and T is the temperature, σ_T is the Thomson cross section, m_e is the electron rest mass, and c is the speed of light. The SZ signal integrated over the cluster extent is proportional to the integrated Compton parameter, such that $D_A^2 Y_{SZ} = (\sigma_T/m_e c^2) \int P dV$, where D_A is the angular distance to the source.

Clusters are currently thought to form via the hierarchical gravitational collapse of dark matter haloes, so that their number as a function of mass and redshift is a sensitive indicator of the underlying cosmology. The ICM is formed when gas falls into the dark matter gravitational potential and is heated to X-ray

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emitting temperatures by shocks and compression. The scale-free nature of this process implies that simple power law relationships exist between the total halo mass and various other physical properties (e.g., [Bertschinger 1985](#); [Kaiser 1986](#)) such as X-ray temperature T or luminosity L (e.g., [Voit 2005](#); [Arnaud et al. 2005, 2007](#); [Pratt et al. 2009](#); [Vikhlinin et al. 2009](#)). As the total mass is not directly observable, such mass proxies are needed to leverage the statistical power of various large-scale surveys for cosmological applications. Since the gas pressure is directly related to the depth of the gravitational potential, the quantity $D_A^2 Y_{SZ}$ is expected to scale particularly closely with the total mass, a claim supported by recent numerical simulations (e.g., [White et al. 2002](#); [da Silva et al. 2004](#); [Motl et al. 2005](#); [Nagai 2006](#); [Wik et al. 2008](#); [Aghanim et al. 2009](#)). SZ surveys for galaxy clusters thus have great potential to produce competitive cosmological constraints.

In a few short years, SZ observations have progressed from the first spatially resolved observations of individual objects ([Pointecouteau et al. 1999, 2001](#); [Komatsu et al. 1999, 2001](#)), to the first discoveries of new objects ([Staniszewski et al. 2009](#)), to large-scale survey projects for cosmology such as the *Atacama Cosmology Telescope* ([Kosowsky 2003](#), *ACT*) and the *South Pole Telescope* ([Carlstrom et al. 2011](#), *SPT*). Indeed, first cosmological results from these surveys have started appearing ([Vanderlinde et al. 2010](#); [Sehgal et al. 2011](#)). Attention is now focussing on the shape and normalisation of the pressure profile (e.g., [Nagai et al. 2007](#); [Arnaud et al. 2010](#); [Komatsu et al. 2011](#)), calibration of the relationship between $D_A^2 Y_{SZ}$ and the total mass for cosmological applications (e.g., [Bonamente et al. 2008](#); [Marrone et al. 2009](#); [Arnaud et al. 2010](#); [Melin et al. 2011](#)), comparison of the measured SZ signal to X-ray predictions ([Lieu et al. 2006](#); [Bielby & Shanks 2007](#); [Afshordi et al. 2007](#); [Komatsu et al. 2011](#); [Melin et al. 2011](#)), and the relationship between $D_A^2 Y_{SZ}$ and its X-ray analogue $Y_{X,500}$ (e.g., [Andersson et al. 2010](#)). First introduced by [Kravtsov et al. \(2006\)](#), the latter quantity is defined as the product of M_g , the gas mass, and T_X , the spectroscopic temperature excluding the core regions. As the link between $Y_{X,500}$ and $D_A^2 Y_{SZ}$ depends on the relationship between the gas mass weighted and X-ray spectroscopic temperatures, it is a sensitive probe of cluster astrophysics.

In the following, we use a subsample of *Planck*¹ Early Release Compact Source Catalogue SZ (ESZ) clusters, consisting of high signal-to-noise ratio *Planck* detections with deep *XMM-Newton* archive observations, to investigate the local ($z \lesssim 0.5$) SZ scaling relations. Given its all-sky coverage and high sensitivity, *Planck* is uniquely suited to this task, allowing high signal-to-noise ratio detection of many hot, massive systems that do not appear in other SZ surveys due simply to their limited area; correspondingly, the large field of view and collecting power of *XMM-Newton* make it the ideal instrument to observe these objects in X-rays out to a significant fraction of the virial radius. Here we investigate the relationship between SZ quantities and X-ray quantities, making full use of the exceptional quality of both data sets. Two complementary companion papers ([Planck Collaboration 2011f](#) and [Planck Collaboration 2011h](#)) harness the statistical power of the *Planck* survey by analysing the SZ flux–X-ray luminosity and SZ flux–optical richness re-

lations, respectively, using a bin-averaging approach. Two further companion papers present the parent catalogue ([Planck Collaboration 2011d](#)) and *XMM-Newton* validation observations of newly-discovered clusters ([Planck Collaboration 2011e](#)).

In this paper we adopt a Λ CDM cosmology with $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_M = 0.3$ and $\Omega_\Lambda = 0.7$. The factor $E(z) = \sqrt{\Omega_M(1+z)^3 + \Omega_\Lambda}$ is the ratio of the Hubble constant at redshift z to its present day value. The variables M_{500} and R_{500} are the total mass and radius corresponding to a total density contrast $500 \rho_c(z)$, where $\rho_c(z)$ is the critical density of the Universe at the cluster redshift; thus $M_{500} = (4\pi/3) 500 \rho_c(z) R_{500}^3$. The quantity $Y_{X,500}$ is defined as the product of $M_{g,500}$, the gas mass within R_{500} , and T_X , the spectroscopic temperature measured in the $[0.15\text{--}0.75] R_{500}$ aperture. The SZ signal is denoted Y_{500} throughout. This quantity is defined by $D_A^2 Y_{500} \equiv (\sigma_T/m_e c^2) \int P dV$. Here D_A is the angular distance to the system, σ_T is the Thomson cross-section, c the speed of light, m_e the electron rest mass, $P \propto n_e T$ is the pressure (the product of the electron number density and temperature), and the integration is performed over a sphere of radius R_{500} . The quantity $D_A^2 Y_{500}$ is the spherically integrated Compton parameter, and Y_{500} is proportional to the flux of the SZ signal within R_{500} .

2. The ESZ catalogue and the *Planck*-*XMM-Newton* archive subsample

2.1. *Planck* and the ESZ Catalogue

Planck ([Tauber et al. 2010](#); [Planck Collaboration 2011a](#)) is the third generation space mission to measure the anisotropy of the cosmic microwave background (CMB). It observes the sky in nine frequency bands covering 30–857 GHz with high sensitivity and angular resolution from $31'$ to $5'$. The Low Frequency Instrument (LFI; [Mandolesi et al. 2010](#); [Bersanelli et al. 2010](#); [Mennella et al. 2011](#)) covers the 30, 44, and 70 GHz bands with amplifiers cooled to 20 K. The High Frequency Instrument (HFI; [Lamarre et al. 2010](#); [Planck HFI Core Team 2011a](#)) covers the 100, 143, 217, 353, 545, and 857 GHz bands with bolometers cooled to 0.1 K. Polarization is measured in all but the highest two bands ([Leahy et al. 2010](#); [Rosset et al. 2010](#)). A combination of radiative cooling and three mechanical coolers produces the temperatures needed for the detectors and optics ([Planck Collaboration 2011b](#)). Two Data Processing Centres (DPCs) check and calibrate the data and make maps of the sky ([Planck HFI Core Team 2011b](#); [Zacchei et al. 2011](#)). *Planck*'s sensitivity, angular resolution, and frequency coverage make it a powerful instrument for galactic and extragalactic astrophysics as well as cosmology. Early astrophysics results are given in *Planck Collaboration (2011e–u)*.

The basic data set used in the present paper is the *Planck* Early Release Compact Source Catalogue SZ (ESZ) sample, described in detail in [Planck Collaboration \(2011d\)](#). The sample is derived from the highest signal-to-noise ratio detections ($S/N > 6$) in a blind multi-frequency search in the all-sky maps from observations obtained in the first ten months of the *Planck* mission.

2.2. The *Planck*-*XMM-Newton* archive subsample

Cross-correlation of the ESZ subsample with the Meta Catalogue of X-ray Clusters (MCXC; [Piffaretti et al. 2010](#)) produced 158 matches with known X-ray clusters. As shown in [Fig. 1](#), these objects lie at a redshift $z \lesssim 0.5$ and cover approx-

¹ *Planck* (<http://www.esa.int/Planck>) is a project of the European Space Agency (ESA) with instruments provided by two scientific consortia funded by ESA member states (in particular the lead countries France and Italy), with contributions from NASA (USA) and telescope reflectors provided by a collaboration between ESA and a scientific consortium led and funded by Denmark.

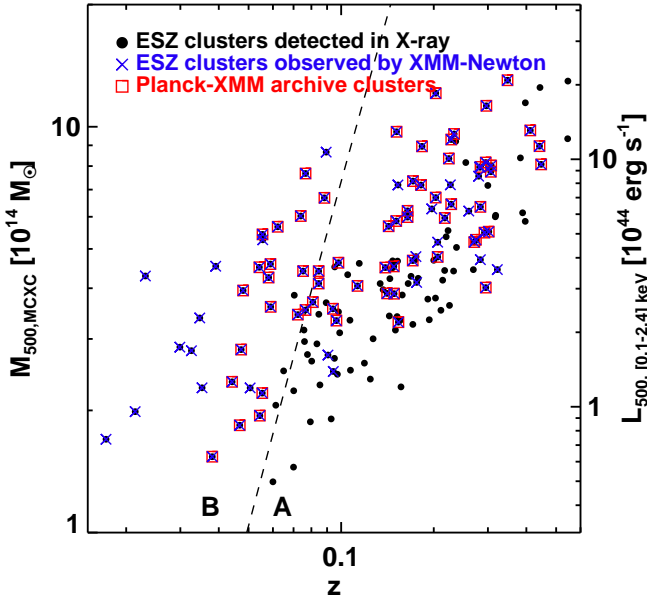


Fig. 1: The 158 *Planck* ESZ clusters already observed in X-rays. Masses are estimated from their X-ray luminosity as described in Piffaretti et al. (2010). The 62 clusters analysed and discussed in this paper are indicated by the red squares. The dashed line represents the locus at which $R_{500} \sim 12'$.

imately a decade in mass². A search for these clusters in the *XMM-Newton* Science Archive³ produced a combined *Planck-XMM-Newton* archive sample of 88 objects as of July 2010, indicated by blue crosses in Fig. 1.

As detailed below in Sect. 3, we used different X-ray data processing techniques depending on cluster angular extent. More specifically, if the source extent lies well within the *XMM-Newton* field of view then the X-ray background can be characterised using a source-free region of the observation, while clusters with a larger angular extent require simultaneous source and background modelling. We label these classes of clusters as A and B, respectively. The dashed line in Fig. 1 illustrates the radius at which $R_{500} \lesssim 12'$, corresponding to the maximum angular extent within which the X-ray background can be characterised in a single *XMM-Newton* field of view. Using this criterion, we divide the *Planck-XMM-Newton* archive sample into 58 A clusters and 30 B clusters.

Not all of the clusters in the full *Planck-XMM-Newton* archive sample are used in the present paper. Some observations in the A cluster list were excluded because soft proton solar flare contamination had rendered the observations unusable, or because the object had not yet been observed at the time of the archive search, or because the target was a clear multiple system unsuited to a spherically-symmetric analysis. For the B clusters, in addition to the high-luminosity systems already published in Bourdin & Mazzotta (2008), we prioritised those where the *XMM-Newton* field of view was expected to cover the largest possible fraction of R_{500} , corresponding to objects with the lowest estimated mass in Fig. 1. The final sample of 62 systems consists of 44 A objects and 18 B objects. While the sample is neither representative nor complete, it represents the largest, highest-quality SZ-X-ray data set currently-available.

² Estimated from the X-ray luminosity-mass relation of Pratt et al. (2009), as detailed in Piffaretti et al. (2010).

³ <http://xmm.esac.esa.int/xsa/>

3. X-ray cluster properties

3.1. X-ray data processing: A clusters

For the A clusters, we use the latest version (v10.0) of the *XMM-Newton*-SAS, ensuring that the most recent calibration corrections are applied to the X-ray photons. Event lists are processed, cleaned for periods of high background, PATTERN-selected and corrected for vignetting as detailed in Pratt et al. (2007).

The quiescent *XMM-Newton* background is dominated by events due to charged particles. We subtract this component by using a background file built from stacked observations obtained with the filter wheel in the CLOSED position, recast to the source position and renormalised using the count rate in the high energy band free of cluster emission⁴. The remaining components are due to the cosmic X-ray background produced by unresolved sources and a diffuse soft X-ray contribution attributable to the Galaxy. The A sample is selected so that $R_{500} \lesssim 12'$, allowing us to model these remaining components using emission from an annular region external to the cluster emission as detailed in Croston et al. (2008) and Pratt et al. (2010).

Point sources were identified from the small scales of wavelet-decomposed images in the [0.3–2] and [2–5] keV bands. These sources were excluded, with the exclusion radius matched to the variation of the PSF size across the detector. We also masked any well-defined substructures that produced prominent secondary maxima and were visible in the larger scales of the wavelet decomposition process.

Surface brightness profiles were extracted from the data in $3' \times 3'$ bins centred on the X-ray peak. Finally, a non-parametric regularisation method was used to derive the deprojected, PSF-corrected density profiles, $n_e(r)$, as described in Croston et al. (2008).

3.2. X-ray data processing: B clusters

For each object in the B cluster sample, a merged energy-position photon cube was built from the various observations of a given target. The cube was built from soft proton-cleaned events from each camera, generated with v10.0 of the *XMM-Newton*-SAS, to which an effective exposure and a background noise array were added. The exposure array was computed from the effective exposure time, with corrections for spatially variable mirror effective areas, filter transmissions, CCD pixel area, chip gaps and bad pixels, as appropriate. The background noise array was modelled as the sum of components accounting for the Galactic foreground and cosmic X-ray background, plus charged particle-induced and out-of-time events. Full details of the method are given in Bourdin & Mazzotta (2008).

The Galactic foreground is a critical model component in the case of the B clusters. These objects often extend over the full *XMM-Newton* field of view, so that the cluster emission cannot be spatially separated from the foreground components. We thus constrained the foreground components using a joint fit of cluster emissivity and temperature in an external annulus corresponding to $\sim R_{500}$. Despite the degeneracy of this estimate with the cluster emissivity itself, in all cases the temperature obtained in this annulus was found to be lower than the average cluster temperature, as is commonly observed in clusters allowing full cluster-foreground spatial separation (e.g., Pratt et al. 2007; Leccardi & Molendi 2008).

⁴ We excluded a $5'$ region around the cluster centre to avoid contamination from residual cluster emission.

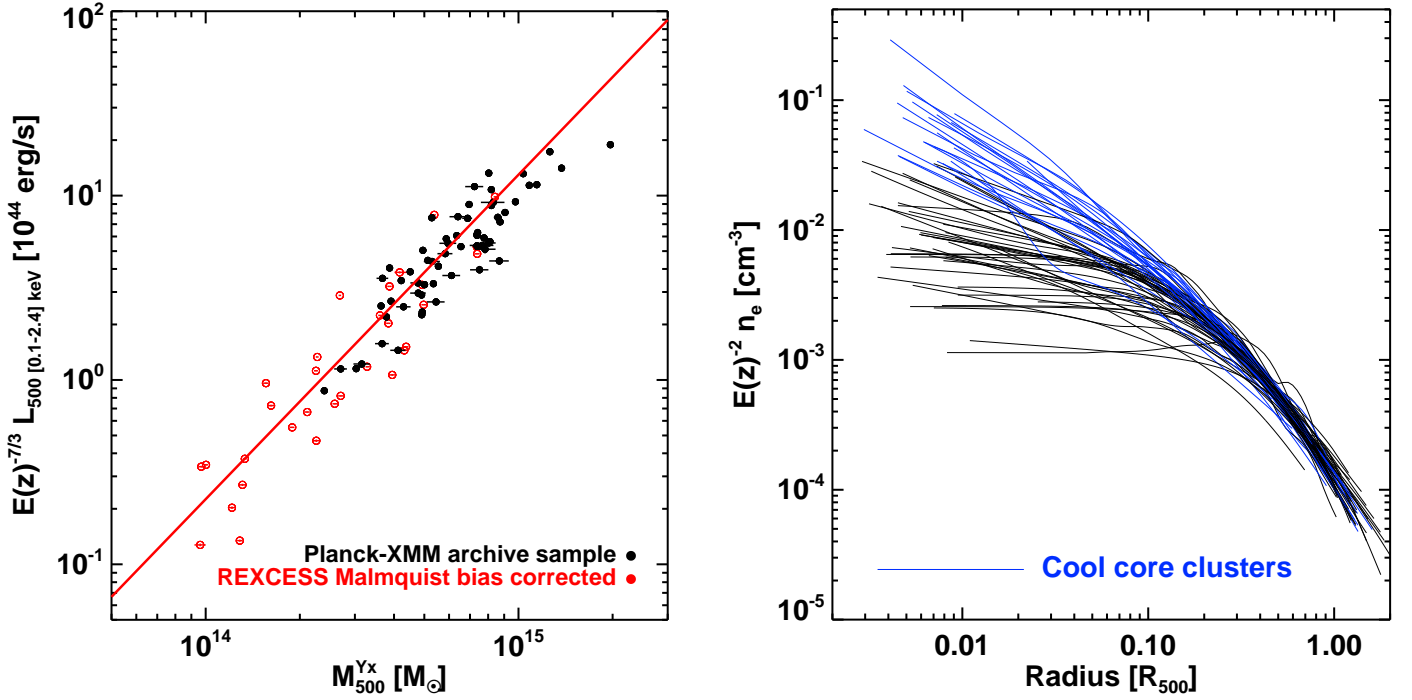


Fig. 2: *Left*: $L_{X,500} - M_{500}$ relation of the *Planck-XMM-Newton* archive sample compared to REXCESS, a representative X-ray cluster sample. Luminosity is estimated interior to R_{500} in the [0.1–2.4] keV band, and mass from the $M_{500} - Y_{X,500}$ relation of [Arnaud et al. \(2010\)](#). The solid red line is the fit to the REXCESS sample only ([Pratt et al. 2009](#)). *Right*: Scaled density profiles of the 62 systems in the *Planck-XMM-Newton* archive sample. Profiles have been corrected for projection and PSF effects as described in the text. Systems classified as cool cores are indicated by blue lines.

The ICM density profiles in the B cluster sample were then derived using the analytic distributions of ICM density and temperature introduced by [Vikhlinin et al. \(2006\)](#). These parametric distributions were projected along the line of sight, convolved with the *XMM-Newton* PSF, and fitted to the observed projected cluster brightness and temperature profiles. The resulting density profiles, $n_e(r)$, were used to derive X-ray quantities for each cluster as described below.

3.3. X-ray quantities

For the current generation of high-resolution X-ray telescopes, the effective limiting radius for high-quality nearby observations of the type discussed here is R_{500} . Beyond this radius, the effect of the variable background becomes dominant and the uncertainties begin to become difficult to quantify. In addition, as shown by [Evrard et al. \(1996\)](#), R_{500} is also the radius within which clusters are relatively relaxed. We estimate the X-ray quantities for each cluster self-consistently within R_{500} using the $M_{500} - Y_{X,500}$ relation given in [Arnaud et al. \(2010, see also Pratt et al. 2010\)](#), assuming standard evolution, viz.,

$$E(z)^{2/5} M_{500} = 10^{14.567 \pm 0.010} \left[\frac{Y_{X,500}}{2 \times 10^{14} \text{ M}_{\odot} \text{ keV}} \right]^{0.561 \pm 0.018} \text{ M}_{\odot}. \quad (1)$$

The radius R_{500} was calculated iteratively as described in [Kravtsov et al. \(2006\)](#). Using Eq. 1 and the definition of M_{500} and $Y_{X,500}$, an equation of the form $R_{500}^3 = C[M_{g,500} T_X]^a$ must be solved. Starting from an initial temperature measurement, the equation is solved for R_{500} , with the gas mass computed from the density profiles discussed above in Sects. 3.1 and 3.2. A new temperature is then estimated within $[0.15 - 0.75] R_{500}$ and the procedure is repeated until convergence. Once converged,

we measure the temperature within an aperture corresponding to R_{500} and calculate the [0.1 – 2.4] keV and [0.5 – 2] keV band luminosities as described in [Pratt et al. \(2009\)](#). The resulting X-ray quantities are listed in Table 1.

The left-hand panel of Fig. 2 shows the $L_{X,500} - M_{500}$ relation of the *Planck-XMM-Newton* archive sample, where $L_{X,500}$ is the X-ray luminosity estimated interior to R_{500} in the [0.1–2.4] keV band⁵, and the mass is estimated from the $M_{500} - Y_{X,500}$ relation given in Equation 1. The data are compared to the equivalent relation from REXCESS, a sample designed to be representative of the X-ray cluster population ([Böhringer et al. 2007](#)). One can see that the *Planck-XMM-Newton* archive clusters are all massive, luminous systems, as expected for objects detected in SZ at high S/N by *Planck*. They follow the general trend exhibited by REXCESS ([Pratt et al. 2009](#)), but extend to higher mass and luminosity.

3.4. Scaled gas density profiles and cool core subsample

The scaled gas density profiles of the full sample of 62 clusters are shown in the right-hand panel of Fig. 2, where each profile has been corrected for evolution and scaled to R_{500} . As has been seen in other cluster samples (e.g., [Croston et al. 2008](#)), there is a large amount of scatter in the central regions, extending out to $\sim 0.15 R_{500}$, beyond which the profiles rapidly converge.

⁵ The standard [0.1 – 2.4] keV band self-similar evolution factor is $E(z)^{-2}$. However, the observed scaling relations are steeper than self-similar (see e.g., [Pratt et al. 2009](#)), leading to a dependence of the evolution factor on the quantity to which it is applied. In the present work we use an evolution factor of $E(z)^{-7/3}$, appropriate for bolometric luminosities, for comparison to REXCESS (see also [Maughan 2007](#)).

Table 1: X-ray and SZ properties.

Name	RA	Dec	z	R_{500}	T_X	$M_{g,500}$	$Y_{X,500}$	$D_A^2 Y_{500}$	M_{500}	$L_{X,500}$	CC
	[deg]	[deg]		[kpc]	[keV]	[$10^{14} M_\odot$]	[$10^{14} M_\odot \text{ keV}$]	[10^{-4} Mpc^2]	[$10^{14} M_\odot$]	[$10^{44} \text{ erg s}^{-1}$]	
RXC J0014.3-3022	3.58	-30.38	0.307	1358	7.72±0.25	1.65±0.01	12.73±0.51	1.74±0.21	9.78±0.21	13.35±0.09	...
A85	10.44	-9.37	0.052	1206	5.78±0.22	0.66±0.01	3.84±0.19	0.47±0.05	5.30±0.31	4.65±0.02	✓
RXC J0043.4-2037	10.84	-20.61	0.292	1152	5.82±0.20	0.88±0.01	5.10±0.20	1.40±0.17	5.88±0.14	8.26±0.08	...
A119	14.02	-1.30	0.044	1114	5.40±0.23	0.45±0.01	2.45±0.14	0.27±0.03	4.12±0.23	1.52±0.01	...
RXC J0232.2-4420	38.06	-44.37	0.284	1223	6.41±0.20	1.07±0.01	6.86±0.26	0.86±0.13	6.95±0.15	12.53±0.09	✓
A401	44.73	13.56	0.075	1355	7.26±0.44	1.02±0.04	7.43±0.58	0.83±0.08	7.65±0.67	5.82±0.04	...
RXC J0303.8-7752	46.00	-77.88	0.274	1251	7.88±0.36	0.96±0.02	7.58±0.45	1.09±0.13	7.37±0.25	7.39±0.07	...
A3112	49.51	-44.26	0.070	1062	5.02±0.15	0.40±0.01	2.03±0.07	0.18±0.03	3.67±0.16	3.84±0.02	✓
A3158	55.72	-53.60	0.060	1124	5.00±0.18	0.53±0.01	2.66±0.12	0.35±0.03	4.29±0.23	2.66±0.01	...
A478	63.35	10.45	0.088	1326	6.43±0.19	1.06±0.03	6.81±0.26	0.92±0.08	7.23±0.48	12.33±0.05	✓
A3266	67.83	-61.42	0.059	1354	7.46±0.22	0.96±0.02	7.17±0.30	0.90±0.07	7.51±0.51	4.22±0.01	...
A520	73.55	2.96	0.203	1325	7.74±0.22	1.13±0.01	8.75±0.32	0.99±0.14	8.11±0.16	7.11±0.04	...
RXC J0516.7-5430	79.17	-54.52	0.295	1266	7.11±0.67	1.20±0.06	8.50±1.06	1.29±0.10	7.82±0.60	7.27±0.38	...
RXC J0528.9-3927	82.22	-39.44	0.284	1218	6.04±0.32	1.11±0.02	6.73±0.46	1.18±0.13	6.88±0.25	10.55±0.11	✓
RXC J0532.9-3701	83.23	-37.02	0.275	1190	6.84±0.26	0.85±0.01	5.82±0.28	0.97±0.13	6.35±0.17	8.40±0.07	✓
RXC J0547.6-3152	86.89	-31.90	0.148	1150	6.10±0.14	0.60±0.01	3.63±0.10	0.45±0.07	5.01±0.08	3.89±0.02	...
A3376	90.47	-39.99	0.045	930	3.39±0.09	0.28±0.01	0.94±0.03	0.10±0.02	2.39±0.06	0.92±0.01	...
RXC J0605.8-3518	91.48	-35.29	0.139	1059	4.93±0.11	0.46±0.01	2.29±0.07	0.47±0.06	3.87±0.06	4.74±0.02	✓
RXC J0645.4-5413	101.39	-54.21	0.164	1303	7.26±0.18	1.01±0.01	7.33±0.24	1.09±0.07	7.40±0.14	7.59±0.04	✓
RXC J0658.5-5556	104.63	-55.96	0.296	1527	11.19±0.25	2.08±0.02	23.22±0.64	2.66±0.14	13.73±0.21	20.05±0.10	...
A665	127.75	65.88	0.182	1331	7.64±0.46	1.12±0.03	8.55±0.61	1.09±0.11	8.04±0.37	6.81±0.10	...
A754	137.24	-9.65	0.054	1423	8.93±0.24	1.04±0.03	9.28±0.39	0.86±0.05	8.69±0.63	4.68±0.02	...
A773	139.49	51.69	0.217	1228	6.78±0.16	0.89±0.01	6.01±0.18	0.86±0.11	6.55±0.11	6.80±0.04	...
A781	140.09	30.49	0.298	1114	5.72±0.10	0.76±0.01	4.32±0.10	0.72±0.14	5.35±0.07	4.75±0.03	...
A868	146.36	-8.64	0.153	1058	4.63±0.16	0.51±0.01	2.34±0.08	0.41±0.07	3.91±0.10	3.18±0.03	...
A963	154.24	39.01	0.206	1123	5.49±0.11	0.66±0.01	3.63±0.09	0.41±0.09	4.95±0.07	6.40±0.03	✓
RXC J1131.9-1955	173.00	-19.92	0.308	1300	7.75±0.31	1.30±0.02	10.11±0.53	1.30±0.23	8.59±0.26	11.01±0.09	...
A1413	178.81	23.39	0.143	1242	6.57±0.07	0.82±0.01	5.41±0.05	0.74±0.08	6.27±0.04	6.85±0.02	✓
RXC J1206.2-0848	181.59	-8.81	0.441	1334	10.15±0.32	1.59±0.02	16.13±0.63	1.70±0.30	10.83±0.24	19.65±0.12	✓
ZwCl1215	184.41	3.65	0.077	1211	6.45±0.27	0.63±0.02	4.09±0.21	0.46±0.07	5.45±0.35	2.88±0.01	...
A1576	189.23	63.19	0.302	1145	6.32±0.47	0.80±0.03	5.05±0.49	0.79±0.11	5.83±0.32	6.94±0.18	...
A3528S	193.65	-29.21	0.053	966	4.11±0.21	0.28±0.01	1.16±0.07	0.22±0.03	2.70±0.13	1.22±0.01	✓
A1644	194.30	-17.40	0.047	1070	4.86±0.20	0.41±0.01	1.99±0.11	0.25±0.04	3.66±0.19	1.66±0.05	✓
A3532	194.39	-30.41	0.056	1015	4.44±0.30	0.34±0.01	1.53±0.12	0.21±0.04	3.16±0.19	1.30±0.01	...
A1650	194.67	-1.76	0.084	1110	5.11±0.06	0.51±0.01	2.61±0.04	0.44±0.06	4.22±0.03	3.79±0.01	✓
A1651	194.88	-4.20	0.084	1135	5.23±0.12	0.56±0.01	2.94±0.08	0.36±0.06	4.51±0.07	4.23±0.02	...
A1689	197.88	-1.35	0.183	1339	8.17±0.12	1.08±0.01	8.84±0.15	1.37±0.15	8.19±0.08	13.29±0.03	✓
A3558	202.00	-31.51	0.047	1170	4.78±0.13	0.67±0.02	3.21±0.13	0.42±0.05	4.77±0.26	3.54±0.01	...
A1763	203.80	41.00	0.223	1275	6.55±0.17	1.14±0.01	7.44±0.27	1.28±0.12	7.37±0.15	8.00±0.05	...
A1795	207.24	26.58	0.062	1254	6.60±0.21	0.73±0.02	4.79±0.18	0.46±0.04	5.96±0.37	5.90±0.02	✓
A1914	216.49	37.83	0.171	1345	8.26±0.19	1.07±0.01	8.80±0.26	1.00±0.09	8.19±0.13	10.73±0.05	...
A2034	227.53	33.49	0.151	1330	7.01±0.15	1.13±0.01	7.94±0.23	0.74±0.10	7.76±0.13	6.99±0.04	...
A2029	227.73	5.75	0.078	1392	7.70±0.41	1.12±0.05	8.63±0.60	0.81±0.07	8.30±0.72	10.00±0.05	✓
A2065	230.61	27.70	0.072	1160	5.36±0.20	0.60±0.02	3.24±0.15	0.39±0.05	4.78±0.28	3.20±0.02	✓
A2163	243.95	-6.13	0.203	1781	13.40±0.45	3.17±0.04	42.51±1.82	4.55±0.21	19.68±0.48	23.86±0.15	...
A2204	248.18	5.59	0.152	1345	7.75±0.21	1.09±0.02	8.45±0.28	1.11±0.10	8.04±0.15	15.73±0.06	✓
A2218	248.99	66.21	0.171	1151	5.23±0.10	0.73±0.01	3.82±0.10	0.77±0.06	5.13±0.08	5.41±0.03	...
A2219	250.10	46.71	0.228	1473	9.37±0.22	1.74±0.02	16.33±0.47	2.34±0.14	11.44±0.20	14.94±0.10	...
A2256	256.13	78.63	0.058	1265	6.40±0.25	0.78±0.02	4.98±0.23	0.71±0.04	6.11±0.40	3.92±0.02	...
A2255	258.24	64.05	0.081	1169	5.79±0.15	0.59±0.01	3.42±0.11	0.52±0.04	4.91±0.09	2.47±0.02	...
RXC J1720.1+2638	260.03	26.61	0.164	1165	5.78±0.12	0.70±0.01	4.02±0.10	0.63±0.08	5.28±0.08	9.14±0.04	✓
A2261	260.61	32.14	0.224	1216	6.23±0.55	0.93±0.04	5.79±0.61	1.18±0.12	6.41±0.41	9.97±0.27	✓
A2390	328.41	17.69	0.231	1423	8.89±0.24	1.54±0.02	13.68±0.46	1.66±0.13	10.35±0.20	17.20±0.09	✓
A3827	330.46	-59.95	0.099	1210	6.19±0.10	0.69±0.01	4.28±0.09	0.63±0.05	5.55±0.07	4.62±0.02	...
RXC J2217.7-3543	334.46	-35.73	0.149	1034	4.68±0.10	0.44±0.01	2.05±0.05	0.35±0.06	3.64±0.06	2.98±0.01	...
RXC J2218.6-3853	334.68	-38.89	0.141	1147	6.19±0.19	0.57±0.01	3.51±0.13	0.34±0.06	4.92±0.11	2.74±0.02	...
RXC J2228.6+2036	337.12	20.62	0.412	1256	8.16±0.30	1.33±0.02	10.86±0.52	1.34±0.23	8.73±0.24	11.96±0.10	...
RXC J2234.5-3744	338.62	-37.75	0.151	1307	7.34±0.12	0.99±0.01	7.24±0.15	0.90±0.07	7.37±0.09	7.21±0.05	...
MACSJ2243.3-0935	340.84	-9.58	0.444	1256	7.98±0.12	1.47±0.01	11.75±0.22	1.91±0.24	9.06±0.10	14.05±0.05	...
A3911	341.60	-52.72	0.097	1066	4.52±0.06	0.48±0.01	2.16±0.04	0.38±0.04	3.78±0.03	2.45±0.01	...
AS1063	342.21	-44.53	0.347	1456	10.73±0.25	1.89±0.02	20.33±0.58	2.21±0.16	12.60±0.20	26.32±0.13	...
A3921	342.49	-64.42	0.094	1071	5.01±0.07	0.44±0.01	2.20±0.02	0.36±0.03	3.82±0.03	2.57±0.02	...

Notes. The temperature T_X is measured in the $[0.15 - 0.75] R_{500}$ region, and the luminosity $L_{X,500}$ is measured interior to R_{500} in the $[0.1 - 2.4]$ keV band. The mass M_{500} is estimated from the $M_{500} - Y_{X,500}$ relation given in Eqn. 1. The final column indicates whether the cluster is classified as a cool core system, defined as described in Sect. 3.4.

It is well-known that some clusters exhibit so-called cool cores, central regions of very dense gas where the cooling time is less than the Hubble time (e.g., Jones & Forman 1984). Such objects have very high X-ray luminosities and extremely low cen-

tral entropies that tend to set them apart from the rest of the X-ray cluster population (e.g., Fabian et al. 1994; Pratt et al. 2010). In addition, the current consensus is that these systems represent a generally more relaxed subset of the cluster population (although

see Burns et al. 2008 for a dissenting view). Following Pratt et al. (2009), we estimated the central gas density $n_{e,0}$ using a β model fit to the gas density profile interior to $0.05 R_{500}$, and classified objects with $E(z)^{-2} n_{e,0} > 4 \times 10^{-2} \text{ cm}^{-3}$ as cool core systems. In total, 22/62 clusters in the present sample are classified as such. These are plotted in blue in Fig. 2 and in all following plots.

4. SZ cluster properties

4.1. Optimisation of the SZ flux extraction

The basic SZ signal extraction procedure is described in full in Planck Collaboration (2011d). In brief, this procedure consists of applying multi-frequency matched filters (MMF, Melin et al. 2006), that incorporate prior knowledge of the signal, to the *Planck* maps. Specifically, the ICM pressure is assumed to follow the universal profile shape derived from the REXCESS sample by Arnaud et al. (2010). The SZ flux is computed by integrating along the line-of-sight and normalising the universal pressure profile. Each profile is truncated at $5 \times R_{500}$, effectively giving a measure of the flux within a cylinder of aperture radius $5 \times R_{500}$, and then converted to the value in a sphere of radius R_{500} for direct comparison with the X-ray prediction. This is the fundamental SZ quantity used in the present paper⁶, and we refer to it throughout as Y_{500} .

Section 6 of Planck Collaboration (2011d) shows that the cluster flux derived from blind application of the MMF algorithm is systematically larger than X-ray expectations. This discrepancy is a result of overestimation of the cluster size θ_{500} due to the freedom to optimise significance with position and size. As shown in Planck Collaboration (2011d), if the SZ signal is instead extracted from a region centred on the X-ray position with size θ_{500} estimated from the X-ray luminosity-mass relation, the SZ flux is in better agreement with X-ray expectations. When additional constraints on the cluster size are available, the SZ flux extraction can be further optimised.

With the present cluster sample we can make use of the higher-quality estimate of the X-ray size θ_{500} , derived from R_{500} , measured using the $M_{500} - Y_{X,500}$ relation as detailed in Sect. 3.3. Appendix A details the improvement in SZ flux extraction when these higher-quality size estimates are used. For each cluster in the sample, we thus re-ran the SZ flux extraction, calculating Y_{500} with the X-ray position and size fixed to the refined values derived from the high-quality *XMM-Newton* observation.

4.2. Robustness tests specific to local sample

Section 6 of Planck Collaboration (2011d) details various robustness tests relevant to all *Planck* SZ papers, including investigation of the cluster size–flux degeneracy discussed above, the impact of the assumed pressure profile used for cluster detection, beam-shape effects, colour corrections, contamination by point sources, and discussion of the overall error budget. For the present sample we undertake two further robustness tests: the first is related to the impact of radio source contamination; the second examines the impact of the assumed pressure profile shape on the derived Y_{500} .

⁶ Note that Y_{500} is the directly observed ‘apparent’ quantity, while $D_A^2 Y_{500}$ is the corresponding ‘absolute’ quantity, intrinsic to the cluster.

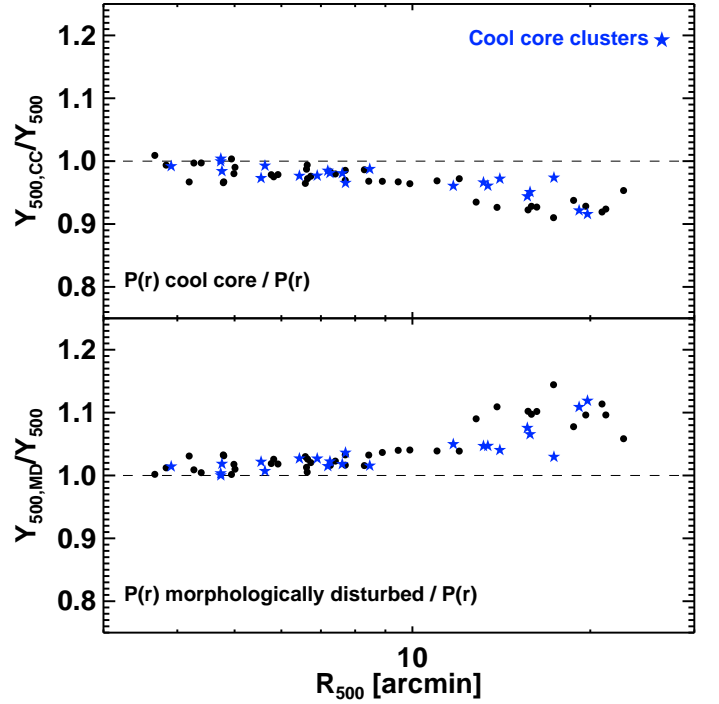


Fig. 3: Comparison of Y_{500} from extraction using the baseline pressure profile with that from the best fitting REXCESS cool core and morphologically disturbed pressure profiles given in Appendix C of Arnaud et al. (2010). Cool core systems are marked as blue stars, other systems as black dots. *Top*: universal pressure profile vs cool-core pressure profile; *Bottom*: universal pressure profile vs morphologically-disturbed pressure profile. The trend with R_{500} is due to the inability of the *Planck* beam to resolve different profiles at small angular size. The effect is small (maximum ~ 10 per cent) and quasi-symmetric (as expected), so no bias is introduced.

4.2.1. Contamination by point sources

Contamination by point sources can affect extraction of the SZ parameters, and have implications for astrophysical studies of clusters or further cosmological applications (Aghanim et al. 2005; Douspis et al. 2006). We have thus checked the possible effect of radio galaxies on the derived Y_{500} by combining data from SUMSS (Bock et al. 1999, a catalogue of radio sources at 0.85 GHz), NVSS (Condon et al. 1998, a catalogue of radio sources at 1.4 GHz), and data from the *Planck* LFI and HFI. Two clusters in our sample exhibit relatively bright ($S(1.4 \text{ GHz}) \gtrsim 1 \text{ Jy}$), flat spectrum radio sources within a radius of $15'$ from the X-ray peak. These sources are clearly seen in LFI data and could potentially affect the SZ measurement. However, as we discuss below in Sect. 5, inclusion or exclusion of these objects has a negligible effect on the derived scaling relations.

4.2.2. Impact of assumptions on pressure profile and scaling

The blind SZ signal detection method used to detect and extract the ESZ clusters from the *Planck* survey (Planck Collaboration 2011d) implements the universal pressure profile from Arnaud et al. (2010). More specifically, the baseline model makes use of the generalised NFW profile fit to the 31 individual REXCESS cluster pressure profiles, after removal of the mass dependence by scaling according to the $M_{500} - Y_{X,500}$ relation given in Eqn. 1. However, Arnaud et al. showed that the scatter of the individ-

ual cluster pressure profiles about the universal form increases toward the central regions, since cool core systems are more peaked, and morphologically disturbed systems are shallower, respectively, than the mean. In their Appendix C, [Arnaud et al.](#) gave the best fitting GNFW model parameters for the average scaled profiles of the REXCESS cool core and morphologically disturbed subsamples.

As our cluster sample contains both cool core and morphologically disturbed systems, it is pertinent to investigate the effect of the baseline pressure profile assumption on the resulting Y_{500} values. We thus re-ran the Y_{500} extraction process separately for each object using the cool core and morphologically disturbed cluster profiles given in Appendix C of [Arnaud et al. \(2010\)](#). The X-ray size θ_{500} is kept the same in each case so that we are investigating the impact of the pressure profile shape within a fixed aperture. Figure 3 shows the ratio of the Y_{500} of the cool core and morphologically disturbed profile extractions to that of the baseline model.

Clear trends are seen in both cases: the ratio tends to increase (decrease) with θ_{500} if the morphologically disturbed (cool core) profile is used instead of the baseline universal profile. Up to $\theta_{500} \sim 10'$ the ratio differs from unity only by 2 per cent on average. Beyond $\theta_{500} \sim 10'$, the derived Y_{500} starts to differ gradually from the baseline value. This effect can be traced to the influence of the *Planck* angular resolution. Since the SZ signal extraction uses all *Planck*-HFI channels, the effective angular resolution is that of the channel with the largest FWHM (~ 10 arcmin at 100 GHz). Below this angular scale the profile shape is washed out by the convolution with the *Planck* beam, while above it, clusters are increasingly well-resolved. The two panels of Fig. 3 show that at the largest θ_{500} the maximum excursion is ~ 10 per cent. Beyond 10 arcmin, the average excursions are ~ 6 and ~ 7 per cent, respectively, for cool-core and morphologically disturbed profiles. Note that the effect is symmetric, in that for large θ_{500} a cool core profile and a morphologically disturbed profile return a value of Y_{500} that differs from the baseline value by approximately the same amount, but the former is lower and the latter is higher.

In the following, the difference in Y_{500} derived from extraction with the cool core and morphologically disturbed cluster profiles is added in quadrature to the uncertainty on the Y_{500} from the baseline extraction. We expect this conservative error estimate to account for any difference in the underlying pressure profile shape from the universal baseline model. As detailed below in Sect. 5, we have further checked the effect of the pressure profile assumption on the derived scaling relation fits, finding it to be entirely negligible.

5. SZ scaling relations

We fitted the parameters governing a scaling relation between $D_A^2 Y_{500}$, the spherically-integrated SZ signal within R_{500} , and its X-ray analogue $Y_{X,500}$. We also fitted parameters governing scaling relations between $D_A^2 Y_{500}$ and various other X-ray-derived quantities including $M_{g,500}$, T_X and $L_{X,500}$. We further investigated the relation between $D_A^2 Y_{500}$ and the total mass, M_{500} , using the $M_{500} - Y_{X,500}$ calibration given in Eqn. 1.

5.1. Fitting method

For each set of observables (B, A) , we fitted a power law relation of the form $E(z)^\gamma D_A^2 Y_{500} = 10^A [E(z)^\kappa X/X_0]^B$, where $E(z)$ is the Hubble constant normalised to its present day value and γ and

κ were fixed to their expected self-similar scalings with z . The fit was undertaken using linear regression in the log-log plane, taking the uncertainties in both variables into account, and the scatter was computed as described in [Pratt et al. \(2009\)](#). In brief, assuming a relation of the form $Y = aX + b$, and a sample of N data points (Y_i, X_i) with errors σ_{Y_i} and σ_{X_i} , the raw scatter was estimated using the error-weighted distances to the regression line:

$$\sigma_{\text{raw}}^2 = \frac{1}{N-2} \sum_{i=1}^N w_i (Y_i - aX_i - b)^2 \quad (2)$$

where

$$w_i = \frac{1/\sigma_i^2}{(1/N) \sum_{i=1}^N 1/\sigma_i^2} \quad \text{and} \quad \sigma_i^2 = \sigma_{Y_i}^2 + a^2 \sigma_{X_i}^2. \quad (3)$$

The intrinsic scatter σ_i was computed from the quadratic difference between the raw scatter and that expected from the statistical uncertainties.

We use the BCES regression method ([Akritas & Bershady 1996](#)), which takes into account measurement errors in both coordinates and intrinsic scatter in the data and is widely used in astronomical regression, giving results that may easily be compared with other data sets fitted using the same method. We fitted all relations using orthogonal BCES regression.

5.2. Effect of point sources and choice of baseline pressure profile

We fitted the SZ scaling relations excluding the two objects with significant radio source contamination (Sect. 4.2.1). For all relations the change in normalisation, slope and intrinsic scatter is negligible compared to their associated uncertainties. We therefore consider the contamination by radio sources to have a negligible effect on scaling relation fits and proceed with the full sample of 62 clusters.

We have also checked whether the best fitting scaling relations are affected by the choice of baseline pressure profile, as follows. For the cool core subsample, we assigned the Y_{500} derived from extraction using the cool core pressure profile. Of the remaining 40 systems, for the 20 objects with the lowest central density (Fig. 2), we assigned the Y_{500} value derived from extraction using the morphologically disturbed profile. We then re-fitted all the scaling relations. The resulting best fits are in full agreement with those derived from the baseline universal profile; i.e., the difference in best fitting parameters (2 per cent maximum) is again smaller than their respective uncertainties.

5.3. Correction for selection bias

It is well known that scaling-relation determinations of the sort we are considering can be biased by selection effects of Malmquist and Eddington type when a significant part of the sample lies near a selection cut (for discussions in a cluster context see e.g., [Vikhlinin et al. 2009](#); [Pratt et al. 2009](#); [Mantz et al. 2010](#) and [Andersson et al. 2010](#)). We estimated the effect of the *Planck* SZ selection as follows. In order to impose a selection cut on the mock catalogues, we used the *observed* relation between $D_A^2 Y_{500}$ and S/N from the region significantly above the selection cut and extrapolated below it, along with an estimate of scatter again from observations, carried out in several redshift bins. We then constructed large mock catalogues of clusters

Table 2: Best fitting orthogonal BCES parameters for scaling relations.

Relation	A_{obs}	B_{obs}	$\sigma_{log,i}$	A_{corr}	B_{corr}	γ	κ	X_0
$D_A^2 Y_{500} - c^* Y_{X,500}$	-4.02 ± 0.01	0.95 ± 0.04	0.10 ± 0.01	-4.02	0.96	–	–	$1 \times 10^{-4} \text{ Mpc}^2$
$D_A^2 Y_{500} - M_{500}$	-4.19 ± 0.01	1.74 ± 0.08	0.10 ± 0.01	-4.22	1.74	$-2/3$	–	$6 \times 10^{14} M_\odot$
$D_A^2 Y_{500} - T_X$	-4.27 ± 0.02	2.82 ± 0.18	0.14 ± 0.02	-4.22	2.92	-1	–	6 keV
$D_A^2 Y_{500} - M_{g,500}$	-4.05 ± 0.01	1.39 ± 0.06	0.08 ± 0.01	-4.03	1.48	$-2/3$	–	$1 \times 10^{14} M_\odot$
$D_A^2 Y_{500} - L[0.1 - 2.4]_{X,500}$	-4.00 ± 0.02	1.02 ± 0.07	0.14 ± 0.02	-3.97	1.12	$-2/3$	$-7/3$	$7 \times 10^{44} \text{ erg s}^{-1}$
$D_A^2 Y_{500} - L[0.5 - 2]_{X,500}$	-3.79 ± 0.03	1.02 ± 0.07	0.14 ± 0.02	-3.75	1.12	$-2/3$	$-7/3$	$7 \times 10^{44} \text{ erg s}^{-1}$

Notes. $c^* = (\sigma_T/m_e c^2)/(\mu_e m_p)$.

Relations are expressed as $E(z)^\gamma [D_A^2 Y_{500}] = 10^A [E(z)^\kappa X/X_0]^B$. The logarithmic intrinsic scatter of the relation is denoted by $\sigma_{log,i}$.

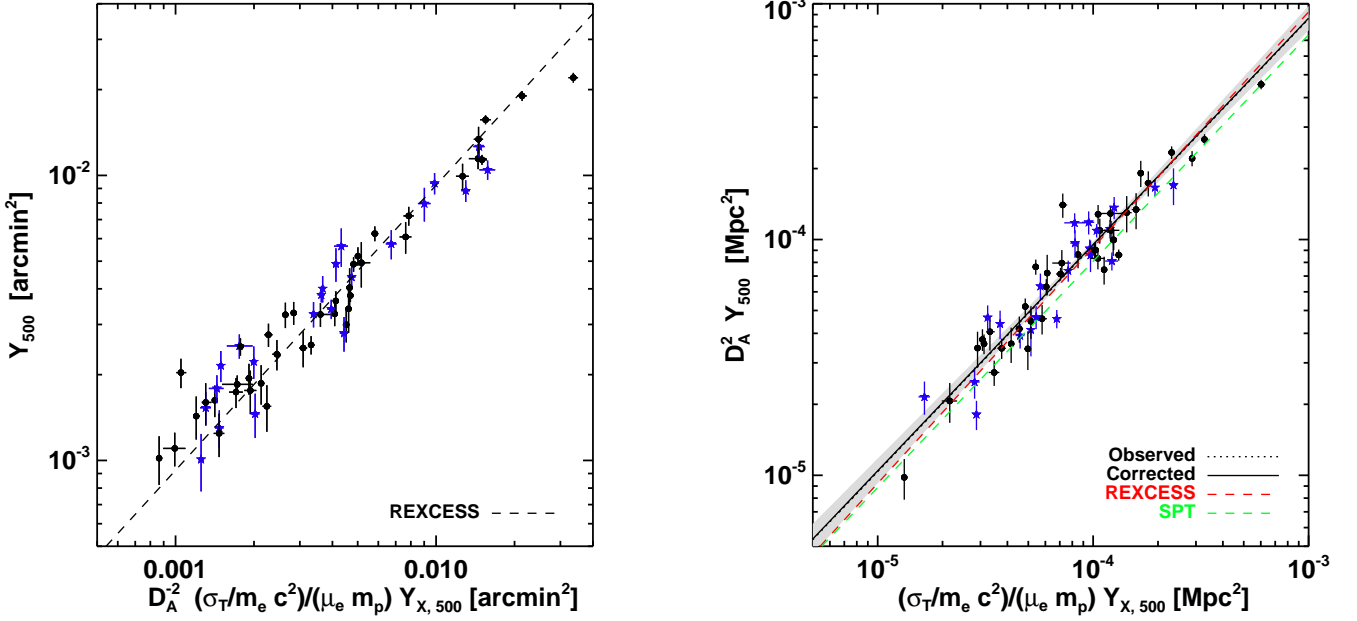


Fig. 4: SZ flux vs X-ray prediction. Blue stars indicate cool core systems. *Left panel*: Relation plotted in units of arcmin^2 . The dashed line is the prediction from REXCESS X-ray observations (Arnaud et al. 2010). *Right panel*: Relation plotted in units of Mpc^2 . The SPT results are taken from Andersson et al. (2010).

through drawing of Poisson samples from a suitably-normalised Jenkins et al. (2001) mass function; to each cluster we assigned a value of $D_A^2 Y_{500}$ by adopting scaling relations with scatter that are consistent with the observed values. This procedure leads to a predicted S/N value that can be used to impose selection cuts on the mock sample. We applied it to the full 158 cluster sample as the only X-ray information needed was the position for SZ signal re-extraction.

The effect on scaling relations is then assessed by assigning further physical properties to the mock catalogue. Following the methods of the X-ray analysis, $Y_{X,500}$ is obtained directly from the mass using Eqn. 1, while $M_{g,500}$ and $L_{X,500}$ are obtained from assumed input scaling relations including scatter. Finally T_X is simply obtained from $Y_{X,500}/M_{g,500}$ on a cluster-by-cluster basis. The input scaling relation slopes and amplitudes are then adjusted until the mock observed samples match those recovered from actual data in Table 2. The input slopes then provide an estimate of the bias-corrected slope that would have been obtained had the bias been absent. The original and bias-corrected estimates are shown in Fig. 4 and 5, and the best-fitting parameters for each relation are given in Table 2. Note that the slopes of the

$Y_{X,500}$ and T_X relations are derived quantities fixed by the other scalings we have chosen.

As seen in Table 2 and in Fig. 5, the importance of the Malmquist correction depends on the relation under consideration. In the case of $Y_{X,500}$ and M_{500} it is negligible, due to the very small scatter seen in these relations. For the other relations, however, the Malmquist corrections can be comparable to the quoted observational uncertainties, indicating that despite the dynamic range of the *Planck* sample, there are biases introduced by the selection cut. The bias-corrected slopes in Table 2 are thus our best current estimates of the true underlying slopes. One should also bear in mind that the bias correction itself carries uncertainty, which we have not been able to estimate, and this increases the uncertainty on the underlying slope.

Note in particular that the bias correction leaves the Y_{500} - $Y_{X,500}$ relation completely consistent with the expected slope of unity, while the relation to M_{500} remains consistent with a slope 5/3.

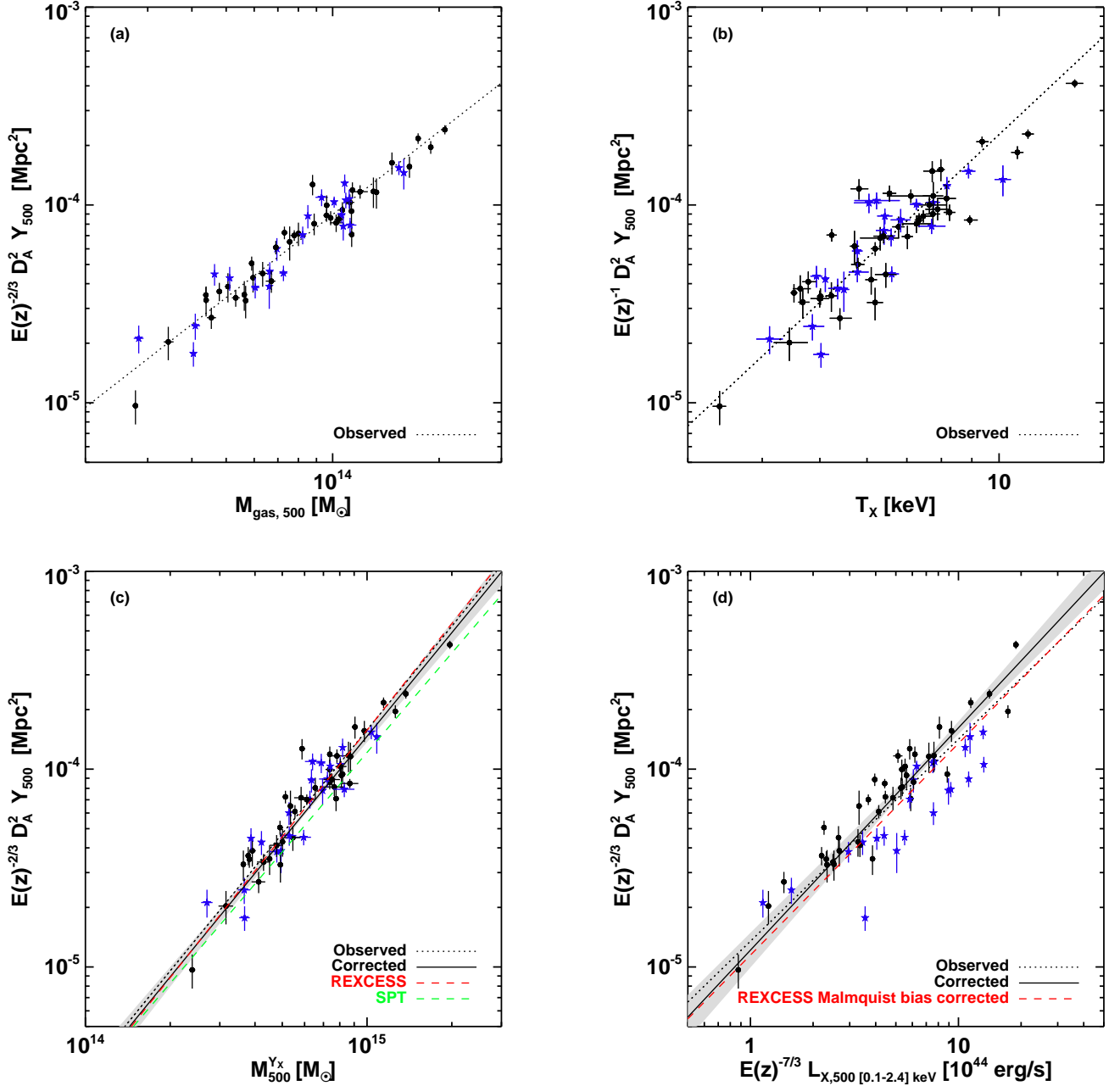


Fig. 5: Scaling relations for the 62 clusters in the *Planck-XMM-Newton* archive sample; fits are given in Table 2. Cool core systems are plotted as blue stars, other systems as black dots. In the upper panels, the dotted line denotes the observed scaling relation fit. In the lower panels, the dotted line denotes the observed scaling relation fit, while the solid line shows the fit once the effects of selection bias are taken into account. The grey shaded area indicates the 1σ uncertainty. The *SPT* results are taken from [Andersson et al. \(2010\)](#).

6. Discussion

6.1. SZ flux vs X-ray prediction

Figure 4 shows the fundamental relation probed by the present study, that between the measured quantities $Y_{X,500}$ and Y_{500} . We recall that the link between these two quantities is sensitive to the structure in temperature and density. Note that X-ray information is used to determine the radius of integration for the SZ signal (i.e., R_{500}) and its overall shape (i.e., the underlying universal pressure profile). However, as we have shown above in Sect. 4.2.2, the amplitude of the SZ signal is relatively insensitive to the assumed pressure profile shape, so that the use of

X-ray priors reduces to a choice of integration aperture. Thus we regard the X-ray and SZ quantities as quasi-independent.

In the left-hand panel of Fig. 4 the relation is plotted in units of arcmin^2 , and shows the excellent agreement between the observed Y_{500} , $Y_{X,500}$ and the X-ray prediction from REXCESS (dashed line). Indeed, fitting the relation with the slope fixed to unity yields a normalisation $Y_{500}/Y_{X,500} = 0.95 \pm 0.03$, perfectly consistent with the value 0.924 ± 0.004 found for the REXCESS sample ([Arnaud et al. 2010](#)), and less than unity as expected for radially-decreasing temperature profiles (e.g., [Pratt et al. 2007](#)). Furthermore, the relation is quite tight (see below), and there is

no indication that cool core systems differ systematically from the other systems.

The right-hand panel of Fig. 4 shows the relation between $Y_{X,500}$ and the spherically-integrated Compton parameter $D_A^2 Y_{500}$. Note in particular that the slope of the bias-corrected relation is completely consistent with unity, and that the intrinsic scatter (calculated as described above in Eqns. 2 and 3) is exceptionally small, $\sigma_{\log,i} = 0.09 \pm 0.01$. A comparison with recent results obtained by *SPT* shows a slight difference in normalisation, although it is not significant given the larger uncertainties in the latter measurement (Andersson et al. 2010).

6.2. Scaling relations

In this Section we investigate other relations between $D_A^2 Y_{500}$ and X-ray quantities. Note that since M_{500} is derived from $Y_{X,500}$, its dependence on $D_A^2 Y_{500}$ is directly linked to the $D_A^2 Y_{500} - Y_{X,500}$ relation with the exception of differing $E(z)$ dependencies. Moreover, $M_{g,500}$ and T_X are not independent; they are related via Eqn. 1. It is still useful to investigate these relations, though, both for completeness and for comparison to recent results from ground-based studies.

Relations between $D_A^2 Y_{500}$ and gas mass $M_{g,500}$ and the X-ray temperature T_X are shown in the upper panels of Fig. 5. The bias-corrected relations yield slopes that are consistent with self-similar (5/3 and 5/2, respectively) to high accuracy. Scatter about the $D_A^2 Y_{500} - M_{g,500}$ relation is small, at $\sigma_{\log,i} = 0.09 \pm 0.01$, while that about the $D_A^2 Y_{500} - T_X$ relation is among the largest of the relations, at $\sigma_{\log,i} = 0.14 \pm 0.02$. Once again, cool core systems are fully consistent with the other clusters, and there is no particular evidence that cool core systems show less scatter than the sample as a whole.

The bottom left-hand panel of Fig. 5 shows the relation between $D_A^2 Y_{500}$ and mass. Here again, the slope of the bias-corrected relation is fully consistent with self-similar (5/3). The scatter is small ($\sigma_{\log,i} = 0.10 \pm 0.01$), although it is a lower limit since the scatter between $Y_{X,500}$ and total mass is not taken into account in deriving M_{500} (it is in fact the same as that about the $D_A^2 Y_{500} - Y_{X,500}$ relation except for the different $E(z)$ scaling). Both slope and normalisation are in excellent agreement with X-ray predictions from REXCESS, as expected from the good agreement in the $D_A^2 Y_{500} - Y_{X,500}$ relation. The slight offset in normalisation of the relation found by *SPT* (Andersson et al. 2010) can be explained by the different calibration of the $M_{500} - Y_{X,500}$ relation used in their study; it is not a significant offset given their larger normalisation uncertainties.

The ease of detecting clusters through their X-ray emission makes the X-ray luminosity an important quantity, and its calibration with the SZ signal is imperative for maximising the synergy between the *Planck* all-sky survey and previous all-sky X-ray surveys such as the RASS and the upcoming *eROSITA* survey. The slope of the $D_A^2 Y_{500} - L_{X,500}$ relation for the present sample, 1.12 ± 0.08 , is in excellent agreement with the slope predicted from X-ray observations alone (1.07 ± 0.08 , Arnaud et al. 2010 from REXCESS), and the normalisation is also consistent within the uncertainties. The slight offset in the best fitting normalisation for the present sample relative to the REXCESS prediction can be attributed to the relative lack of strong cooling core clusters in the present sample compared to REXCESS (see Fig. 2). The scatter, $\sigma_{\log,i} = 0.14 \pm 0.02$, is largest about this relation due to the influence of cool cores, which are segregated from the other systems and all lie to the high-luminosity side. Indeed, as Fig. 5 shows and Table 2 quantifies, while the vast

majority of the dispersion about the $D_A^2 Y_{500} - L_{X,500}$ relation is due to cool cores, these systems do not contribute significantly to the dispersion about the $D_A^2 Y_{500} - Y_{X,500}$ relation. Thus while the X-ray luminosity is very sensitive to the presence of cool cores, $D_A^2 Y_{500}$ appears to be less so.

The slope of our best fitting $D_A^2 Y_{500} - L_{X,500}$ relation is also fully consistent within 1σ with that derived by Planck Collaboration (2011f), which is based on a bin-averaging analysis at the position of known X-ray clusters in the MCXC (Piffaretti et al. 2010). As X-ray selection is more sensitive to the presence of cool cores (due to the density squared dependence of the X-ray luminosity), one might expect the Planck Collaboration (2011f) best-fitting $D_A^2 Y_{500} - L_{X,500}$ relation to be shifted to slightly higher luminosities (i.e., a slightly lower normalisation), as is seen. However, since the MCXC selection function is both complex and unknown, it is impossible to correct their relation for Malmquist bias effects. Thus some part of the normalisation difference between the two studies arises from correction for selection effects. However the good agreement between our results and those from Planck Collaboration (2011f)’s bin-averaged analysis argues that sample selection does not have a strong effect on the scaling relations derived from the latter analysis.

Finally, the results we have derived show that X-ray and SZ measurements give a fully coherent view of cluster structure out to moderately large scales. In particular, they indicate that effects of clumping in the X-ray gas are not significant, at least in the mass and radial range we have probed in the present study. Furthermore, the excellent agreement between the observed $D_A^2 Y_{500} - Y_{X,500}$ relation and the X-ray predictions argue that the SZ and X-ray calibrations we have used are fundamentally sound.

7. Conclusions

We have presented SZ and X-ray data from a sample of 62 local ($z < 0.5$) galaxy clusters detected at high S/N in the *Planck* survey and observed by *XMM-Newton*. The objects range over approximately a decade in mass ($M_{500} \sim 2 - 20 \times 10^{14} M_\odot$), and, while the sample is neither representative nor complete, it constitutes the largest, highest-quality SZ–X-ray data set currently available. This study has been undertaken in the framework of a series of papers on cluster astrophysics based on the first *Planck* all-sky data set (Planck Collaboration 2011d,e,f,h).

SZ and X-ray quantities have been extracted within R_{500} and we have presented a detailed study of the resulting SZ scaling relations. Moreover, we have investigated how selection effects influence the results of the scaling relation fits. Their influence is subtle, but the slopes and normalisations of the scaling relations are generally in good agreement with X-ray predictions and other results after accounting for the selection effects. For the fundamental $D_A^2 Y_{500} - Y_{X,500}$ relation, we measure a remarkably small logarithmic intrinsic scatter of only (10 ± 1) per cent, consistent with the idea that both quantities are low-scatter mass proxies.

The results are fully consistent with the predictions from X-ray observations (Arnaud et al. 2010) and with recent measurements from a smaller sample spanning a wider redshift range observed with *SPT* (Andersson et al. 2010). The results are also in excellent agreement with the statistical analysis undertaken at the positions of known X-ray clusters (Planck Collaboration 2011f). This excellent agreement between observed SZ quantities and X-ray-based predictions underlines the robustness and

consistency of our overall view of ICM properties. It is difficult to reconcile with the claim, based on a recent *WMAP* 7-year analysis, that X-ray data over-predict the SZ signal (Komatsu et al. 2011).

The results presented here, derived from only 62 systems, provide a maximally-robust local reference for evolution studies or for the use of SZ clusters for cosmology. Overall, the agreement between the present results, ground-based results and X-ray predictions augurs well for our understanding of cluster astrophysics and for the use of clusters for cosmology. Future work will make use of the individual pressure profile shape as derived from X-rays to further improve the SZ flux extraction. Comparison of X-ray and SZ pressure profiles will also be undertaken, as will comparison of measurements with independent mass estimation methods. We will also extend our analysis to the full *Planck* catalogue, observing higher-redshift systems, to constrain evolution, and lower-mass objects, to better probe cluster astrophysics.

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References

- Afshordi, N., Lin, Y., Nagai, D., & Sanderson, A. J. R. 2007, *MNRAS*, 378, 293
- Aghanim, N., da Silva, A. C., & Nunes, N. J. 2009, *A&A*, 496, 637
- Aghanim, N., Hansen, S. H., & Lagache, G. 2005, *A&A*, 439, 901
- Akritas, M. G. & Bershad, M. A. 1996, *ApJ*, 470, 706
- Andersson, K., Benson, B., Ade, P., et al. 2010, *arXiv:1006.3068*
- Arnaud, M., Pointecouteau, E., & Pratt, G. W. 2005, *A&A*, 441, 893
- Arnaud, M., Pointecouteau, E., & Pratt, G. W. 2007, *A&A*, 474, L37
- Arnaud, M., Pratt, G. W., Piffaretti, R., et al. 2010, *A&A*, 517, A92
- Bersanelli, M., Mandolesi, N., Butler, R. C., et al. 2010, *A&A*, 520, A4+
- Bertschinger, E. 1985, *ApJS*, 58, 39
- Bielby, R. M. & Shanks, T. 2007, *MNRAS*, 382, 1196
- Bock, D., Large, M. I., & Sadler, E. M. 1999, *AJ*, 117, 1578
- Böhringer, H., Schuecker, P., Pratt, G. W., et al. 2007, *A&A*, 469, 363
- Bonamente, M., Joy, M., LaRoque, S. J., et al. 2008, *ApJ*, 675, 106
- Bourdin, H. & Mazzotta, P. 2008, *A&A*, 479, 307
- Burns, J. O., Hallman, E. J., Gantner, B., Motl, P. M., & Norman, M. L. 2008, *ApJ*, 675, 1125
- Carlstrom, J. E., Ade, P. A. R., Aird, K. A., et al. 2011, *PASP*, 123, 568
- Condon, J. J., Cotton, W. D., Greisen, E. W., et al. 1998, *AJ*, 115, 1693
- Croston, J. H., Pratt, G. W., Böhringer, H., et al. 2008, *A&A*, 487, 431
- da Silva, A. C., Kay, S. T., Liddle, A. R., & Thomas, P. A. 2004, *MNRAS*, 348, 1401
- Douspis, M., Aghanim, N., & Langer, M. 2006, *A&A*, 456, 819
- Evrard, A. E., Metzler, C. A., & Navarro, J. F. 1996, *ApJ*, 469, 494
- Fabian, A. C., Crawford, C. S., Edge, A. C., & Mushotzky, R. F. 1994, *MNRAS*, 267, 779
- Jenkins, A., Frenk, C. S., White, S. D. M., et al. 2001, *MNRAS*, 321, 372
- Jones, C. & Forman, W. 1984, *ApJ*, 276, 38
- Kaiser, N. 1986, *MNRAS*, 222, 323
- Komatsu, E., Kitayama, T., Suto, Y., et al. 1999, *ApJ*, 516, L1
- Komatsu, E., Matsuo, H., Kitayama, T., et al. 2001, *PASJ*, 53, 57
- Komatsu, E., Smith, K. M., Dunkley, J., et al. 2011, *ApJS*, 192, 18
- Kosowsky, A. 2003, *New A Rev.*, 47, 939
- Kravtsov, A. V., Vikhlinin, A., & Nagai, D. 2006, *ApJ*, 650, 128
- Lamarre, J., Puget, J., Ade, P. A. R., et al. 2010, *A&A*, 520, A9+
- Leahy, J. P., Bersanelli, M., D'Arcangelo, O., et al. 2010, *A&A*, 520, A8+
- Leccardi, A. & Molendi, S. 2008, *A&A*, 486, 359
- Lieu, R., Mittaz, J. P. D., & Zhang, S. 2006, *ApJ*, 648, 176
- Mandolesi, N., Bersanelli, M., Butler, R. C., et al. 2010, *A&A*, 520, A3+
- Mantz, A., Allen, S. W., Ebeling, H., Rapetti, D., & Drlica-Wagner, A. 2010, *MNRAS*, 406, 1773
- Marrone, D. P., Smith, G. P., Richard, J., et al. 2009, *ApJ*, 701, L114
- Maughan, B. J. 2007, *ApJ*, 668, 772
- Melin, J., Bartlett, J. G., & Delabrouille, J. 2006, *A&A*, 459, 341
- Melin, J., Bartlett, J. G., Delabrouille, J., et al. 2011, *A&A*, 525, A139+
- Mennella et al. 2011, Planck early results 03: First assessment of the Low Frequency Instrument in-flight performance (Submitted to *A&A*, [arXiv:astro-ph/1101.2038])
- Motl, P. M., Hallman, E. J., Burns, J. O., & Norman, M. L. 2005, *ApJ*, 623, L63
- Nagai, D. 2006, *ApJ*, 650, 538
- Nagai, D., Kravtsov, A. V., & Vikhlinin, A. 2007, *ApJ*, 668, 1
- Piffaretti, R., Arnaud, M., Pratt, G. W., Pointecouteau, E., & Melin, J. 2010, *arXiv:1007.1916*
- Planck Collaboration. 2011a, Planck early results 01: The Planck mission (Submitted to *A&A*, [arXiv:astro-ph/1101.2022])
- Planck Collaboration. 2011b, Planck early results 02: The thermal performance of Planck (Submitted to *A&A*, [arXiv:astro-ph/1101.2023])
- Planck Collaboration. 2011c, Planck early results 07: The Early Release Compact Source Catalogue (Submitted to *A&A*, [arXiv:astro-ph/1101.2041])
- Planck Collaboration. 2011d, Planck early results 08: The all-sky early Sunyaev-Zeldovich cluster sample (Submitted to *A&A*, [arXiv:astro-ph/1101.2024])
- Planck Collaboration. 2011e, Planck early results 09: XMM-Newton follow-up for validation of Planck cluster candidates (Submitted to *A&A*, [arXiv:astro-ph/1101.2025])
- Planck Collaboration. 2011f, Planck early results 10: Statistical analysis of Sunyaev-Zeldovich scaling relations for X-ray galaxy clusters (Submitted to *A&A*, [arXiv:astro-ph/1101.2043])
- Planck Collaboration. 2011g, Planck early results 11: Calibration of the local galaxy cluster Sunyaev-Zeldovich scaling relations (Submitted to *A&A*, [arXiv:astro-ph/1101.2026])
- Planck Collaboration. 2011h, Planck early results 12: Cluster Sunyaev-Zeldovich optical Scaling relations (Submitted to *A&A*, [arXiv:astro-ph/1101.2027])
- Planck Collaboration. 2011i, Planck early results 13: Statistical properties of extragalactic radio sources in the Planck Early Release Compact Source Catalogue (Submitted to *A&A*, [arXiv:astro-ph/1101.2044])
- Planck Collaboration. 2011j, Planck early results 14: Early Release Compact Source Catalogue validation and extreme radio sources (Submitted to *A&A*, [arXiv:astro-ph/1101.1721])
- Planck Collaboration. 2011k, Planck early results 15: Spectral energy distributions and radio continuum spectra of northern extragalactic radio sources (Submitted to *A&A*, [arXiv:astro-ph/1101.2047])
- Planck Collaboration. 2011l, Planck early results 16: The Planck view of nearby galaxies (Submitted to *A&A*, [arXiv:astro-ph/1101.2045])
- Planck Collaboration. 2011m, Planck early results 17: Origin of the submillimetre excess dust emission in the Magellanic Clouds (Submitted to *A&A*, [arXiv:astro-ph/1101.2046])
- Planck Collaboration. 2011n, Planck early results 18: The power spectrum of cosmic infrared background anisotropies (Submitted to *A&A*, [arXiv:astro-ph/1101.2028])
- Planck Collaboration. 2011o, Planck early results 19: All-sky temperature and dust optical depth from Planck and IRAS — constraints on the “dark gas” in our Galaxy (Submitted to *A&A*, [arXiv:astro-ph/1101.2029])
- Planck Collaboration. 2011p, Planck early results 20: New light on anomalous microwave emission from spinning dust grains (Submitted to *A&A*, [arXiv:astro-ph/1101.2031])
- Planck Collaboration. 2011q, Planck early results 21: Properties of the interstellar medium in the Galactic plane (Submitted to *A&A*, [arXiv:astro-ph/1101.2032])
- Planck Collaboration. 2011r, Planck early results 22: The submillimetre properties of a sample of Galactic cold clumps (Submitted to *A&A*, [arXiv:astro-ph/1101.2034])
- Planck Collaboration. 2011s, Planck early results 23: The Galactic cold core population revealed by the first all-sky survey (Submitted to *A&A*, [arXiv:astro-ph/1101.2035])
- Planck Collaboration. 2011t, Planck early results 24: Dust in the diffuse interstellar medium and the Galactic halo (Submitted to *A&A*, [arXiv:astro-ph/1101.2036])
- Planck Collaboration. 2011u, Planck early results 25: Thermal dust in nearby molecular clouds (Submitted to *A&A*, [arXiv:astro-ph/1101.2037])
- Planck Collaboration. 2011v, The Explanatory Supplement to the Planck Early Release Compact Source Catalogue (ESA)

- Planck HFI Core Team. 2011a, Planck early results 04: First assessment of the High Frequency Instrument in-flight performance (Submitted to A&A, [arXiv:astro-ph/1101.2039])
- Planck HFI Core Team. 2011b, Planck early results 06: The High Frequency Instrument data processing (Submitted to A&A, [arXiv:astro-ph/1101.2048])
- Pointecouteau, E., Giard, M., Benoit, A., et al. 1999, ApJ, 519, L115
- Pointecouteau, E., Giard, M., Benoit, A., et al. 2001, ApJ, 552, 42
- Pratt, G. W., Arnaud, M., Piffaretti, R., et al. 2010, A&A, 511, A85+
- Pratt, G. W., Böhringer, H., Croston, J. H., et al. 2007, A&A, 461, 71
- Pratt, G. W., Croston, J. H., Arnaud, M., & Böhringer, H. 2009, A&A, 498, 361
- Rosset, C., Tristram, M., Ponthieu, N., et al. 2010, A&A, 520, A13+
- Sehgal, N., Trac, H., Acquaviva, V., et al. 2011, ApJ, 732, 44
- Staniszewski, Z., Ade, P. A. R., Aird, K. A., et al. 2009, ApJ, 701, 32
- Sunyaev, R. A. & Zeldovich, Y. B. 1972, Comments on Astrophysics and Space Physics, 4, 173
- Tauber, J. A., Mandolesi, N., Puget, J., et al. 2010, A&A, 520, A1+
- Vanderlinde, K., Crawford, T. M., de Haan, T., et al. 2010, ApJ, 722, 1180
- Vikhlinin, A., Burenin, R. A., Ebeling, H., et al. 2009, ApJ, 692, 1033
- Vikhlinin, A., Kravtsov, A., Forman, W., et al. 2006, ApJ, 640, 691
- Voit, G. M. 2005, Reviews of Modern Physics, 77, 207
- White, M., Hernquist, L., & Springel, V. 2002, ApJ, 579, 16
- Wik, D. R., Sarazin, C. L., Ricker, P. M., & Randall, S. W. 2008, ApJ, 680, 17
- Zacchei et al. 2011, Planck early results 05: The Low Frequency Instrument data processing (Submitted to A&A, [arXiv:astro-ph/1101.2040])

Appendix A: Optimised SZ extraction and comparison with X-ray predictions

As discussed in the main text, with the present cluster sample we have optimised the SZ photometry by using the higher-quality estimate of the X-ray size θ_{500} , derived from R_{500} measured using the $M_{500} - Y_{X,500}$ relation as detailed in Sect. 3.3.

In Fig. A.1 we examine the change in Y_{500} when derived using different characteristic sizes θ_{500} to extract the SZ signal. We also compare the SZ signal predicted using X-ray observations (from the $Y_{500}/Y_{X,500}$ relation of Arnaud et al. 2010) to the observed SZ signal. In all cases, the ICM pressure is assumed to follow the baseline universal profile of Arnaud et al. (2010). As extensively described in Planck Collaboration (2011d), the SZ flux Y_{500} is computed by integrating along the line-of-sight and normalising the universal pressure profile. Each profile is truncated at $5 \times R_{500}$, effectively giving a measure of the flux within a cylinder of aperture radius $5 \times R_{500}$, and then converted to the value in a sphere of radius R_{500} for direct comparison with the X-ray prediction.

The left hand panel (a) of Fig. A.1 shows the relation between Y_{500} and that predicted from *XMM-Newton* observations as used in the present paper, illustrating the tight agreement between the two quantities. We recall that here, the R_{500} within which the SZ signal is extracted is derived from the measured $Y_{X,500}$ using the $M_{500} - Y_{X,500}$ relation given in Eqn. 1.

The middle panel (b) of Fig. A.1 assumes that only the X-ray position and luminosity of the cluster are known. In this case the mass is derived from the $M_{500} - L_{X,500}$ relation as described in the MCXC of Piffaretti et al. (2010), thus yielding the characteristic size used to extract the SZ signal, θ_{500} . This mass is also used to predict $Y_{X,500}$ via the $M_{500} - Y_{X,500}$ relation in Eqn. 1. Consistently, the expected SZ signal is extracted from a region of size θ_{500} centred on the X-ray position given in MCXC (as in Planck Collaboration 2011f).

In the right hand panel (c) of Fig. A.1 the position and size of the cluster are unknown, thus Y_{500} is derived blindly together with the SZ flux. However, the predicted SZ flux is derived as above in panel (b). The agreement between measured and predicted values clearly degrades dramatically from panels (a) to (c). Comparing panels (b) and (a), there is a systematic shift to

lower predicted $Y_{X,500}$ values, with a segregation now appearing between cool cores and the other systems. This can be explained by the fact that using the luminosity as a simple mass proxy leads to an underestimate of the mass for morphologically disturbed systems in view of their position with respect to the mean $L_{X,500} - M_{500}$ relation (Pratt et al. 2009). The inverse effect is seen for the cool cores. In addition, there is a smaller impact on the measured Y_{500} via the effect of the assumed θ_{500} . However the effect is smaller: the average ratio of *XMM-Newton* and MCXC characteristic sizes $\theta_{500,MCXC}/\theta_{500,XMM}$ is 0.95 ± 0.06 , corresponding to a change in area of ~ 10 per cent, which translates into a similar variation in SZ flux. This shows that the X-ray luminosity in the MCXC is a sufficiently good mass proxy for a reliable size estimate.

Finally, panel (c) of Fig. A.1 illustrates the size-flux degeneracy in blind *Planck* measurements. When Y_{500} is measured blindly, the size is on average overestimated (see also Planck Collaboration 2011d), and so the disagreement with predictions is even more apparent.

We see that as a result of the size-flux degeneracy, an accurate estimate of the characteristic size is mandatory in order to derive an accurate measure of Y_{500} . A similar conclusion was reached in Planck Collaboration (2011d, see their Fig. 11), where the effect was demonstrated using the full sample of 158 clusters known in X-rays (i.e., those included in the MCXC). These authors found that, in addition to a reduction in intrinsic scatter (from 43 to 34 per cent), knowledge of the cluster size dramatically reduced the offset of the measured Y_{500} to that predicted from X-rays (from 84 to 14 per cent – compare panels (c) and (b) of Fig. A.1 above).

However as noted in Planck Collaboration (2011d), there is still a small but systematic discrepancy. This is mostly due to the use of $L_{X,500}$, a quantity which shows considerable scatter with mass, as a mass proxy. The superior constraints provided by the *XMM-Newton* observations on the cluster size and on $Y_{X,500}$ suppress most of this remaining systematic effect (compare panels (b) and (a)). A smaller contribution is linked to effects due to the nature of the sample selection. This illustrates that a fully coherent approach is needed when undertaking a proper comparison between SZ and X-ray predictions.

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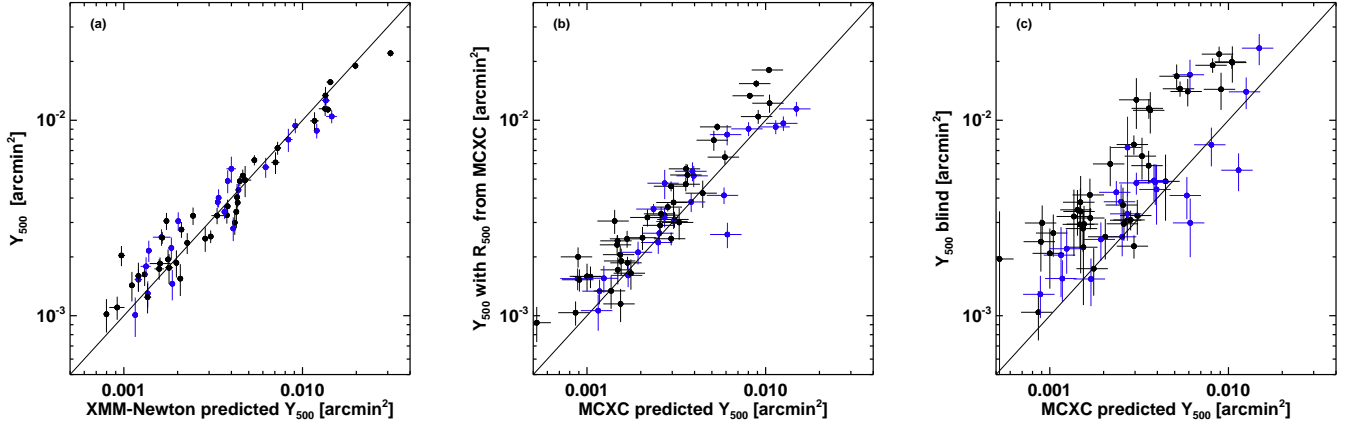


Fig. A.1: Comparison of the *Planck* measured SZ flux, Y_{500} and the predictions from X-ray measurements, $(\sigma_T/m_e c^2)/(\mu_e m_p) D_A^{-2} Y_{X,500}$. Cool core systems are marked as blue stars, other systems as black dots. (Left) $Y_{X,500}$ from *XMM-Newton* (see top-left panel of Fig. 5), and Y_{500} computed at the position and θ_{500} derived from *XMM-Newton* measurements. (Middle) $Y_{X,500}$ and Y_{500} respectively from $L_{X,500}$, and position and θ_{500} as given in the MCXC (Piffaretti et al. 2010). (Right) $Y_{X,500}$ from $L_{X,500}$ as given in the MCXC, while Y_{500} are blind *Planck* measurements.

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